

# The Tully-Fisher and Faber-Jackson Relations

There are three important properties for galaxies: radius  $R$ , luminosity  $L$ , and mass  $M$ , which are connected to “observables” with:

$$\theta = R/d \quad (1)$$

$$f = L/(4\pi d^2) \quad (2)$$

$$v^2 = GM/R \quad (3)$$

where  $\theta$  is the angular diameter,  $d$  is the distance galaxy-observer, and the third relation is the virial theorem. The surface brightness  $I$  is independent of distance and is given by:

$$I = f/\theta^2 = Lv^4/(4\pi G^2 M^2) \quad (4)$$

and then

$$L = v^4/I \times 1/[4\pi G^2 (M/L)^2] \quad (5)$$

where  $M/L$  is the mass-to-light ratio. If **we can assume** that, for a give class of galaxies, the central surface brightness  $I$  and the  $M/L$  are constant, then

$$L \propto v^4 \quad (6)$$

which is the basis of **the most useful distance indicators in cosmology**.

For galaxy disks, where the rotation velocity is measured from HI 21cm line, the observed relation “the Tully-Fisher relation”, is:

$$L \propto V_{rot}^{3-4}, \quad (7)$$

where the precise exponent depend on the band of observations, e.g  $L_B$  or  $L_R$ ...

For elliptical parameters, the appropriate velocity is the central velocity dispersion which can be measured from the size of the absorption lines in the stellar spectrum. The “the Faber-Jackson relation” is observed:

$$L \propto \sigma_0^4. \quad (8)$$

However, this relation is very scattered and there is the need to introduce another parameter.

## The Fundamental Plane

The observational quantities are: effective radius  $r_e$ , luminosity  $L$ , mean surface brightness  $\langle I \rangle_e$  (called also  $\langle \mu_e \rangle$ ), velocity dispersion  $\sigma_V$ .

From the observational point of view,  $\langle I_e \rangle$  correlates with  $L$  and  $\sigma_V$  correlates with  $r_e$ , but a much better correlation is obtained combining  $\sigma_V$ ,  $\langle I_e \rangle$ , and  $r_e$ , i.e.

The origin of the fundamental plane is the following. Assume that ellipticals form a “homologous” family, e.g. dispersion  $\sigma_V$ ; note that  $L$ ,  $\langle I \rangle_e$ ,  $\langle \mu_e \rangle$  are not all independent since they are connected by the typical  $I$  profile of elliptical galaxies (de Vaucouleurs law):

$$L = c_1 I_e r_e^2, \quad (9)$$

Moreover, from the virial theorem

$$M = c_2 \frac{\sigma_0^2 r_e}{G}. \quad (10)$$

Combining them, we obtain

$$r_e = c_2/c_1 \times (M/L)^{-1} \sigma_0^2 I_e^{-1}. \quad (11)$$

If one has  $M/L \propto L^{\sim 0.2}$ , one recovers the observed fundamental plane. Alternatively, one could have constant  $M/L$  with a structure which varies relative to one of more of the fundamental variables. This study of this topic is ongoing.