The Aurora



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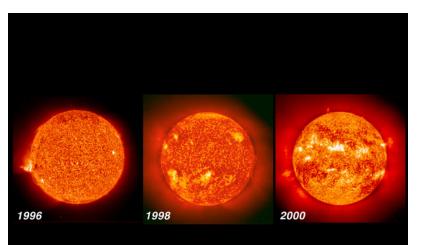


HISTORY

- China: Flying Dragons (2000 aC)
- Bible (Antico Testamento, Ezechiele: Ed ecco un vento tempestoso avanzarsi dal Settentrione, una grande nube che splendeva tutt'intorno, un fuoco in cui guizzavano bagliori, e nel centro come lo splendore dell'eletto in mezzo al fuoco)
- Ancient Greece (Senofane: Cumuli di nubi ardenti)
- Halley: Orientation of Auroral Curtains aligned with Earth's Magnetic Field (1716, Marzo)
- Loomise: Auroral Oval (1859)
- Becqerel: Solar Particles (1878)
- Birkeland: Field Alligned Currents (1902-3)
- Störmer: Motion of Charged Particles in Earth's Magnetic Field, Height of the Aurora (1907)
- Vegard: Proton Aurora (1939)
- Anger: First Global Space-Based Auroral Image (ISIS-2) (inizio anni '70)

SOLAR WIND

Flow of charged particles - *plasma* - streaming from the Sun's *corona* in *all directions*



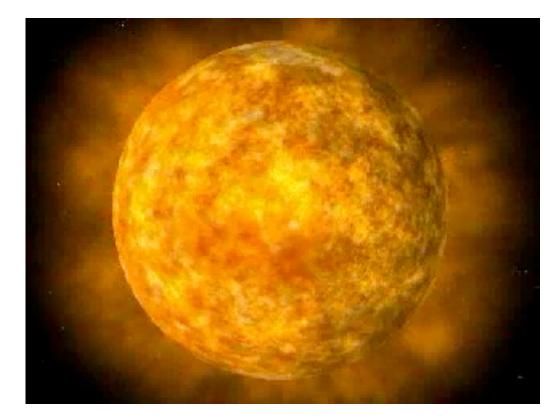
Min	Av	Max
1	3	100
200	400	900
0.4	6.5	100
0	5	25
0.2	6	80
	1 200 0.4 0	132004000.46.505

Comparison of the Solar Corona at Solar Maximum and Minimum (White Light Eclipse Images from the High Altitude Observatory)

February 16,1980 Solar Eclipse (Near Solar Maximum) November 3, 1994 Solar Eclipse (Near Solar Minimum)

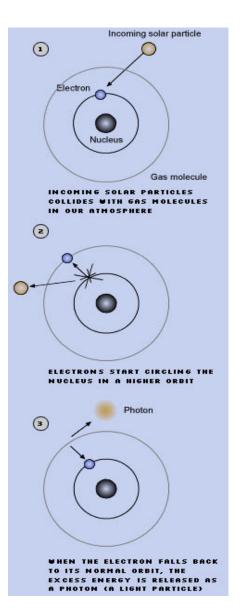
WHAT CAUSES THE AURORA?

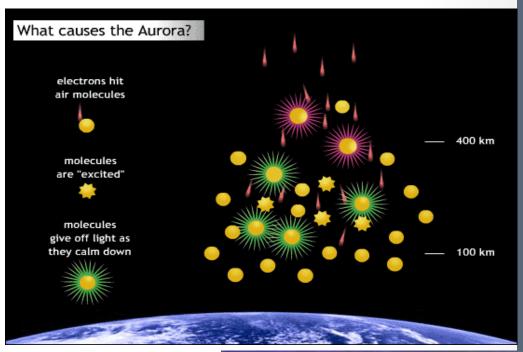
The aurora is a visible manifestation of the solar-terrestrial connection



Solar wind particles, following the geomagnetic field lines, precipitate into the upper atmosphere where interact with atmospheric atoms

WHAT CAUSES THE AURORA?



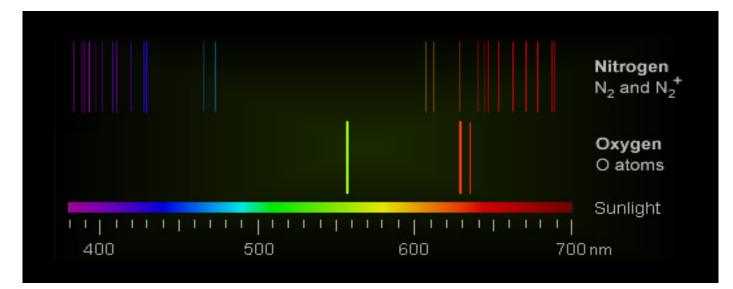


 $hv = \Delta E$



AURORA SPECTRUM

The aurora is a source of X-rays,UV, IR, radio and optical radiation → electron interacts with neutral atoms or molecules which emits radiation due to disexitation



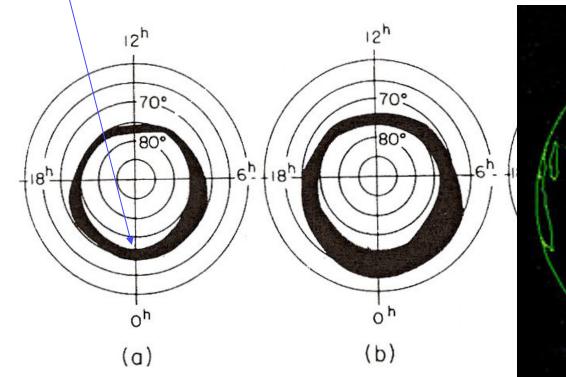
Green colour: atomic oxygen line at 557.7 nm, 100-200 km altitude Red colour: atomic oxygen spectral triplet at 630.0, 636.4, 639.1 nm, h >200 km Violet or Blue colour: molecular nitrogen line at 391.4, 427.0 ,470.0 nm, h <100 km

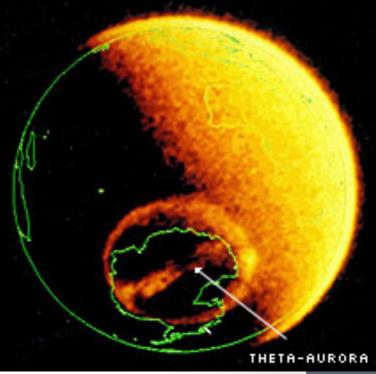
AURORAL OVAL

Seen from space, northen and southern lights appear as oval shaped circle with the magnetic pole in the centre

Stretched out untisunward ~67^o geomagnetic latitude

Compressed sunward ~78^o geomagnetic latitude



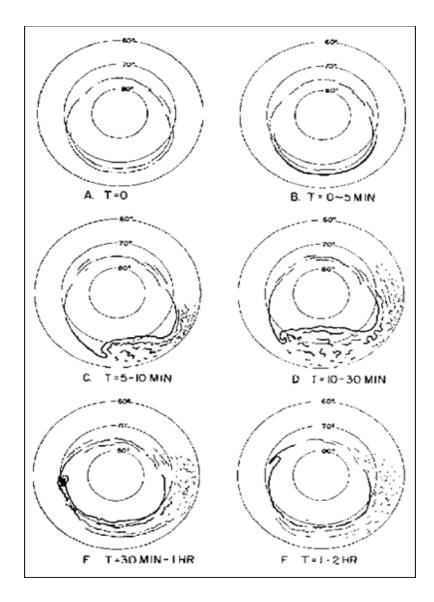


AURORA OCCURRENCE

Within the auroral zone, the aurora can be seen during clear winter night

- 27-day intervals
- more frequent in late autumn and early spring
- Northern lights activity corresponds closely to sunspot activity, which follows an 11-year cycle
- 20-30% less during solar minimum than at solar maximum

THE AURORAL SUBSTORM



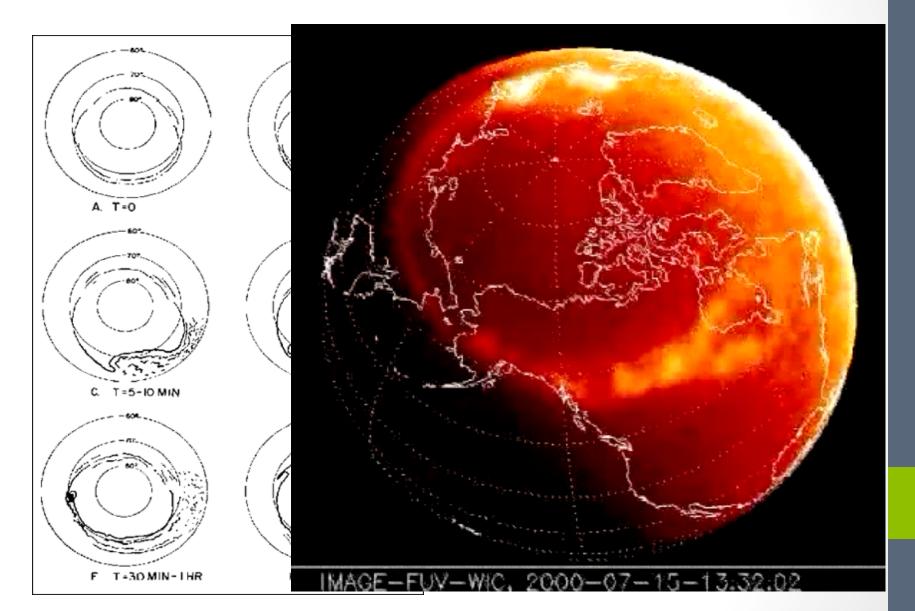
T = 0 min →Quiet phase: weak arcs in the midnight regions at high geomagnetic latitude (~75°)

T = 0~5 min → Oneset phase: southernmost arc brightens and moves southward

T = 5~10 min → Expansion phase: the bright arc forms a bulge wich expands northward and moves westward at high speed

Akasofu, 1968

The Auroral Substorm



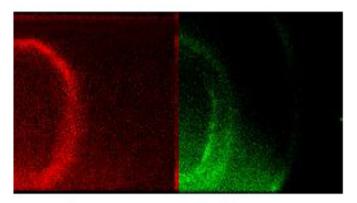
AURORAL STRUCTURE



- Arcs: green, sometimes with red above and purple below, quiet phase
- Bands or Arcs with Structure:
 - green, sometimes with red above and purple below, expansion phase
- Corona: Geometric perspective effect, expansion phase
- Diffuse Glows: greenish at high latitudes, red at lower latitudes, recovery phase, faint

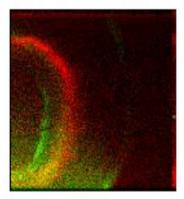
PROTON AURORA

Electron and proton auroras are different and develop differently over time



proton aurora

electron aurora



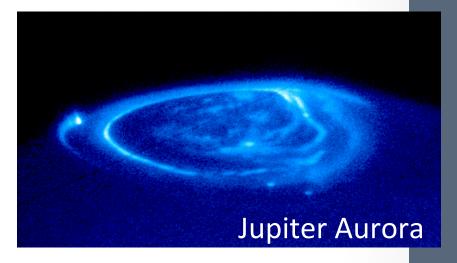
proton + electron aurora

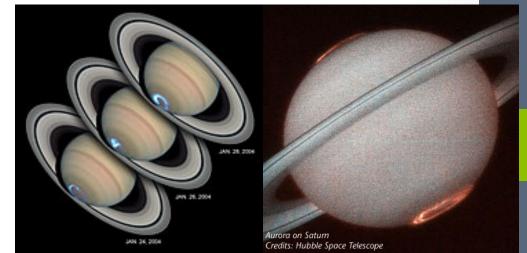
- equally bright but less structured
- Protons quickly become neutralized as they combine with electrons
- very important at the start of the substorm

AURORA ON OTHER PLANETS

Aurora can occur on every planet or moon with a magnetic field and an atmosphere

The process of generating auroras is the same throughout the whole solar system even if the configuration, the colours and particularly the rapidlyvarying displays, are different from what we see on earth





Saturn Aurora

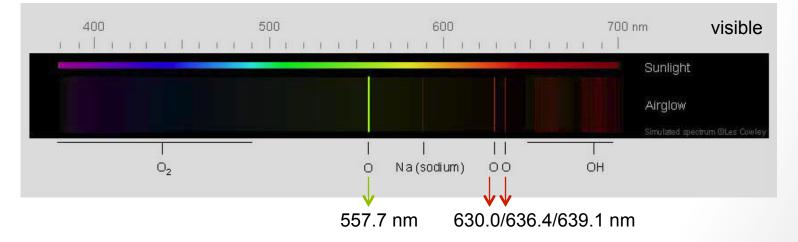
AIRGLOW

Emission of photons from atmospheric constituents excited by the electromagnetic radiation from the sun.

Atomic oxygen lines (100-160km):

 $O + O + M \rightarrow O_2' + M'$ $O_2' + O \rightarrow O_2 + O'$





REFERENCES

• BOOKS

Handbook of the Solar-Terrestrial Environment, Y. Kamide & A.Chian Editors, Springer

Physics of the Upper Polar Atmosphere, Asgeir Brekke, WILEY

• LINKS

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