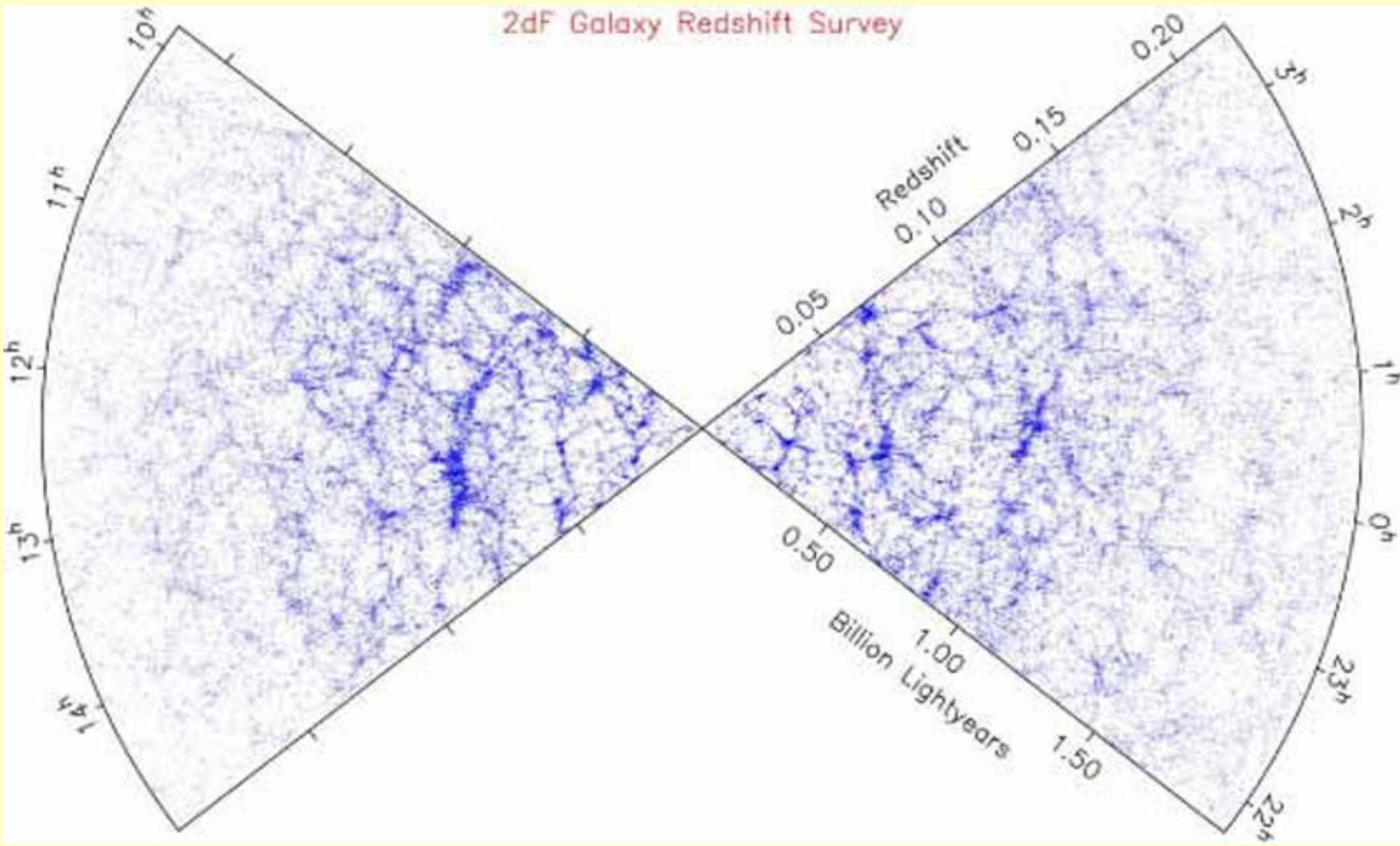
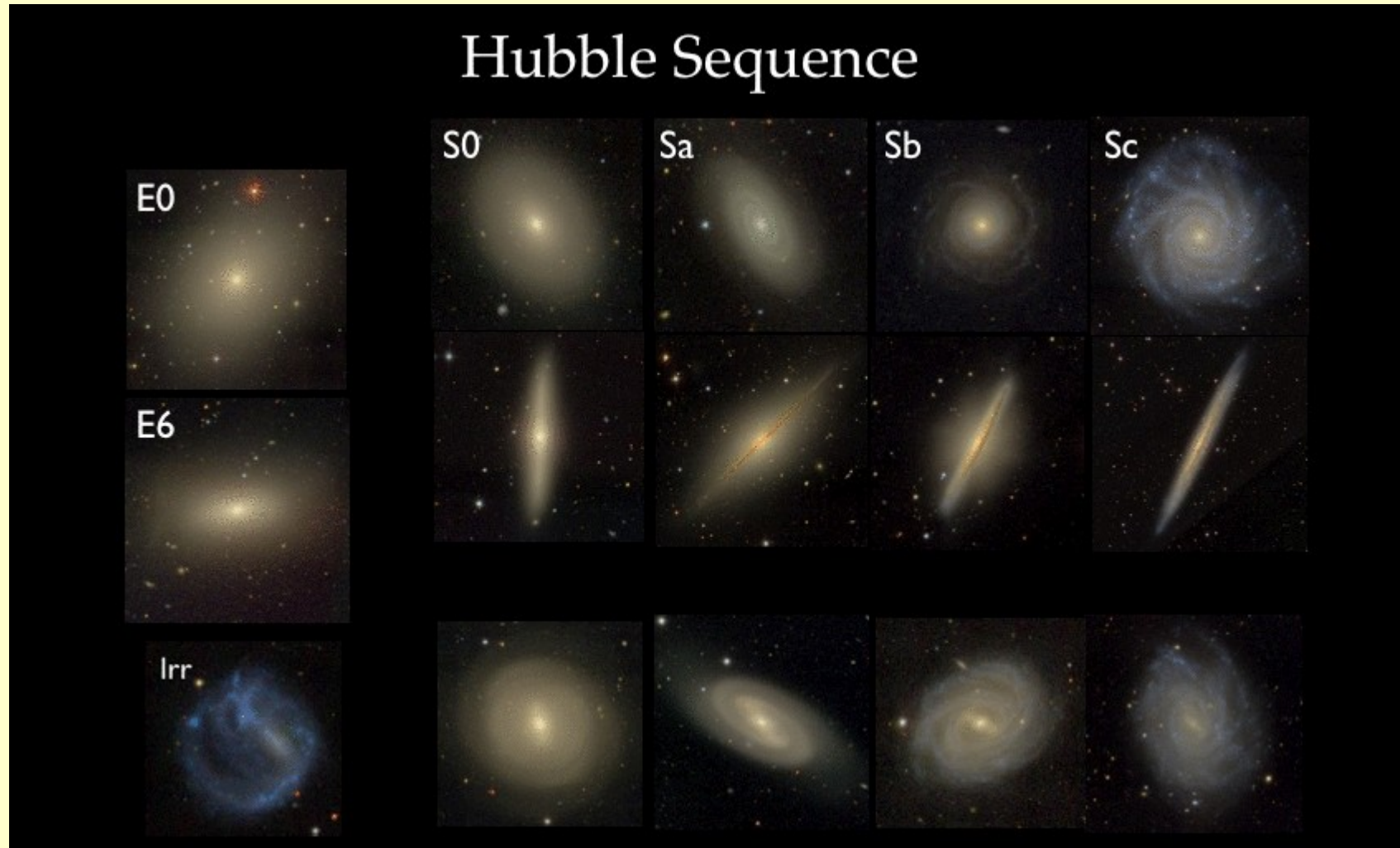


# Galaxy distribution – Colless 2001



Galaxies differ for their morphologies

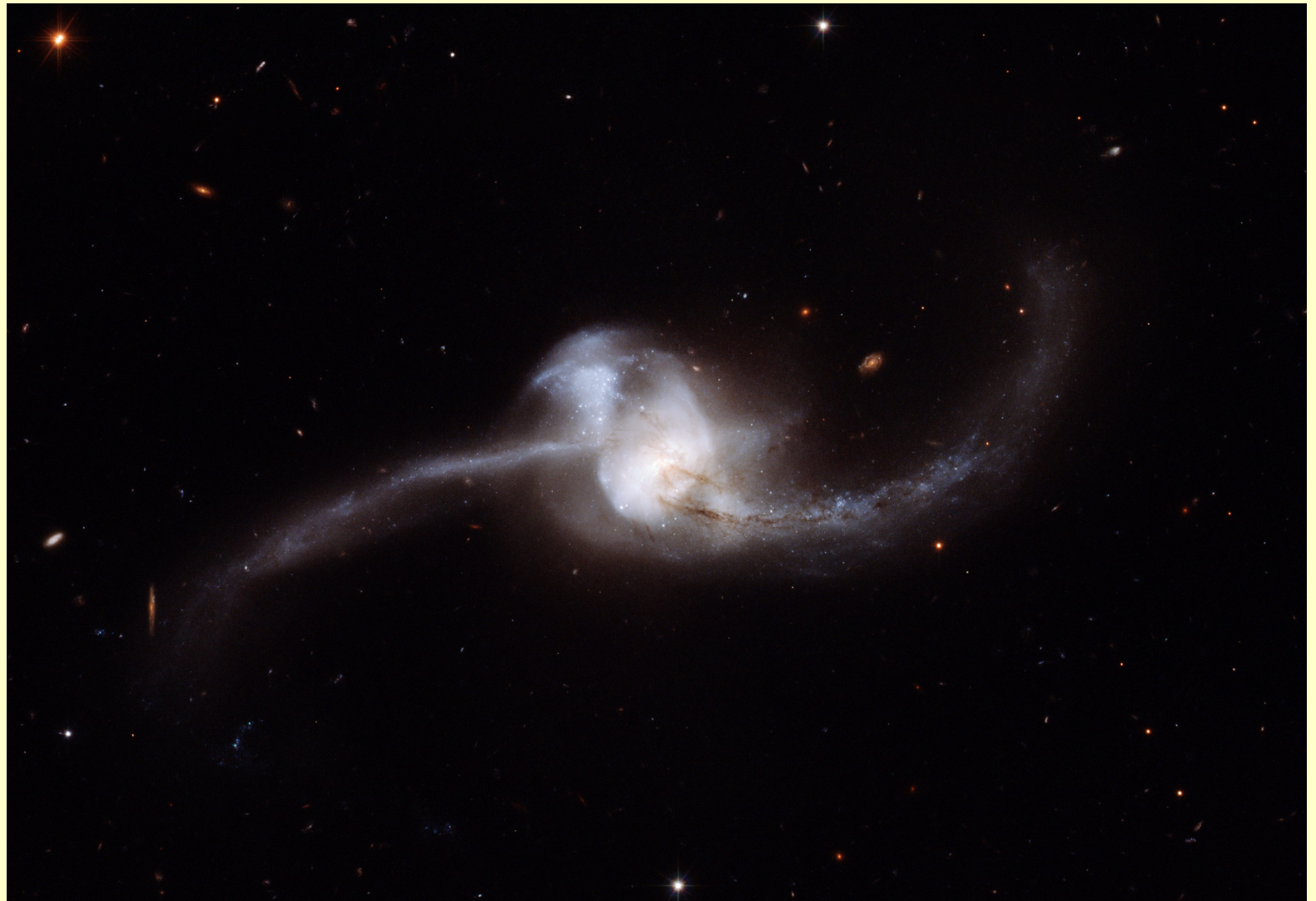


Credit to NASA-HST

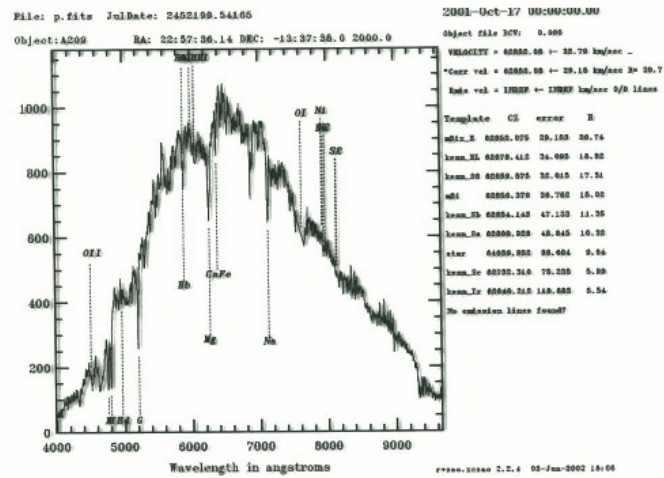
# Merging Galaxies and Environmental Effects

Credit NASA – Martin – here you can see tails of the tidal interaction

Merger: 2 galaxies  $\rightarrow$  1 galaxy

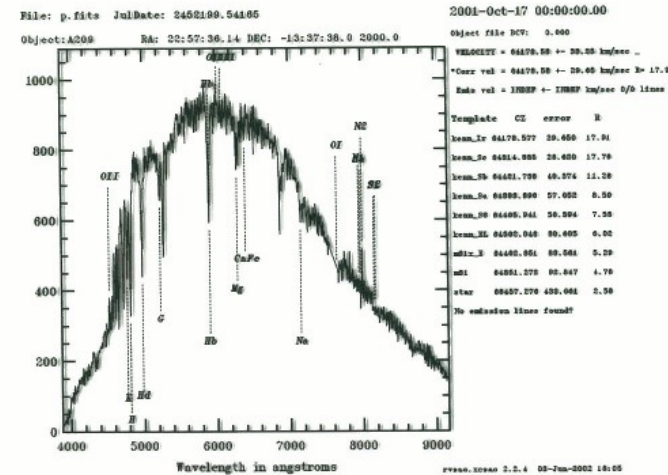


Galaxies differ for their spectra



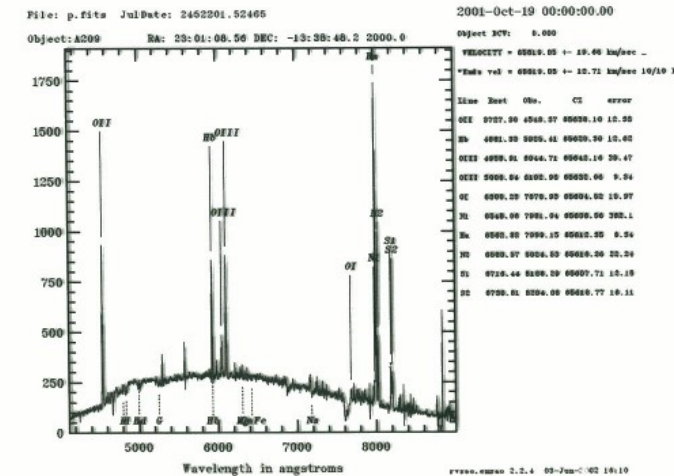
CD GALAXY

deep H $\beta$



E+A GALAXY

POST-STARBUST GALAXY?



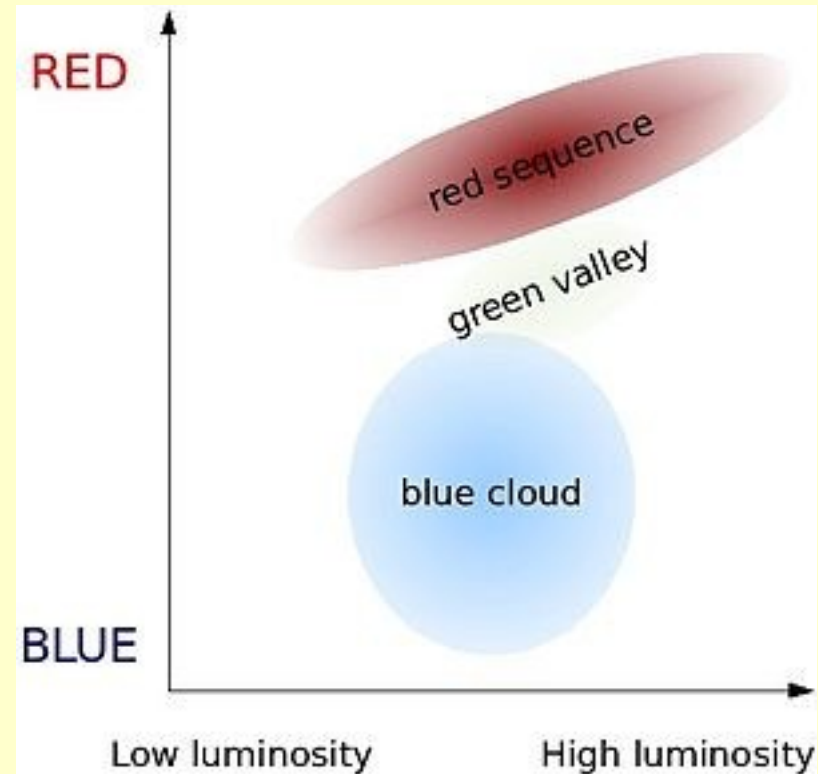
EMISSION LINE GALAXY

... STARBUST?

Abell 209  
Girardi+Mercurio(PhD)  
obs. NTT

Galaxies differ for their colors

Color-magnitude diagram



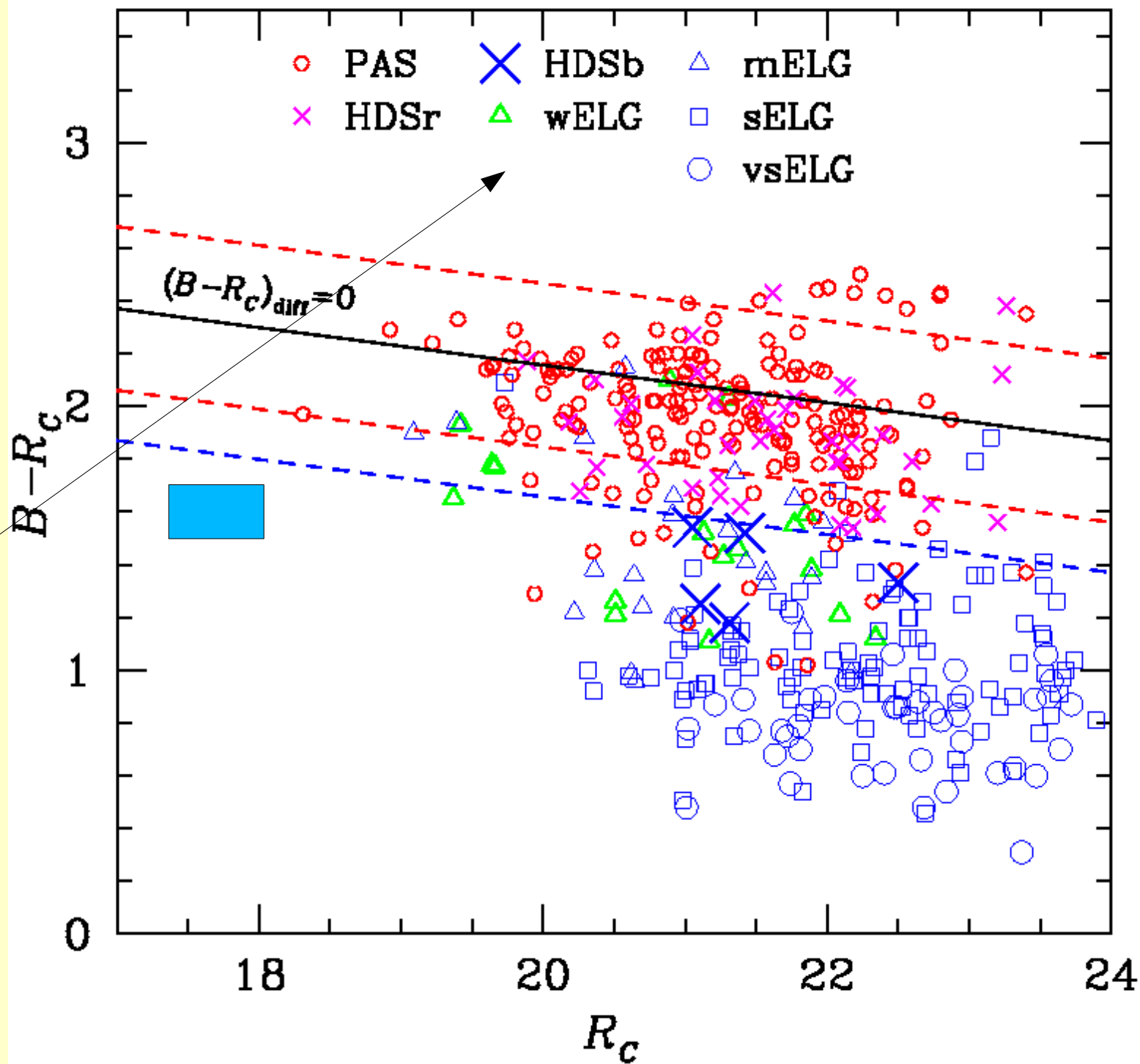
A mock-up of the galaxy color–magnitude diagram with three populations: the red sequence, the blue cloud, and the green valley. Credit to Wikipedia website

Galaxy dichotomy: EARLY TYPES Red Passive Spirals/Irregulars Galaxies  
LATE TYPES Blue Star-Forming Ellipticals Galaxies

Color-Mag  
Relation for  
Cluster galaxies  
in  
MACSJ1206

Girardi+2015

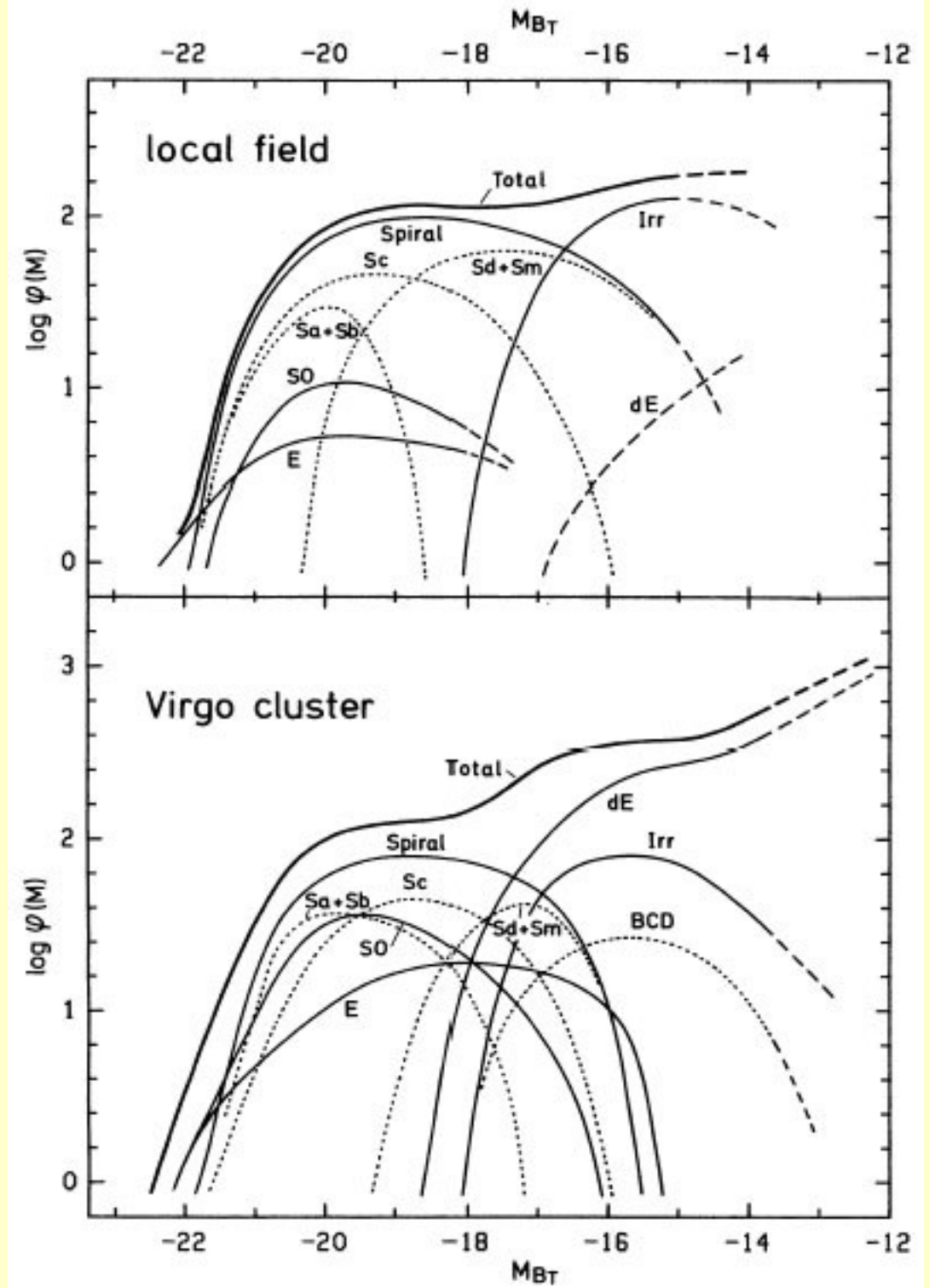
Different  
spectral types



Field and clusters  
differ for their morphological  
Content as shown by their  
LF (Many more Ellipticals  
in clusters!)

Kraan-Korteweg & Tamman 1979

Kraan-Korteweg & Binggeli 1987



Early galaxies populate more clusters than field and more central/dense regions than external regions.

## Morphology density relation

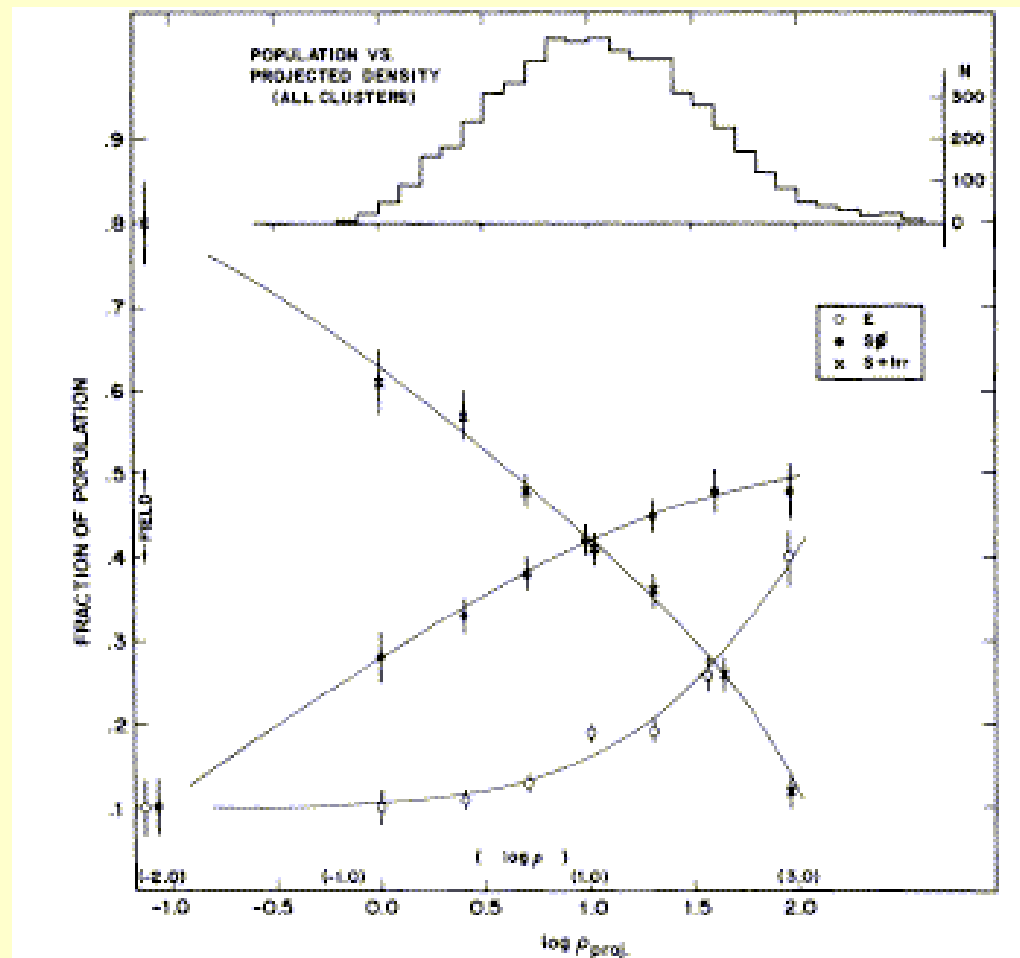


FIG. 4.—The fraction of E, S0, and S+I galaxies as a function of the log of the projected density, in galaxies  $\text{Mpc}^{-2}$ . The data shown are for all cluster galaxies in the sample and for the field. Also shown is an estimated scale of true space density in galaxies  $\text{Mpc}^{-2}$ . The upper histogram shows the number distribution of the galaxies over the bins of projected density.



# Morphology density relation

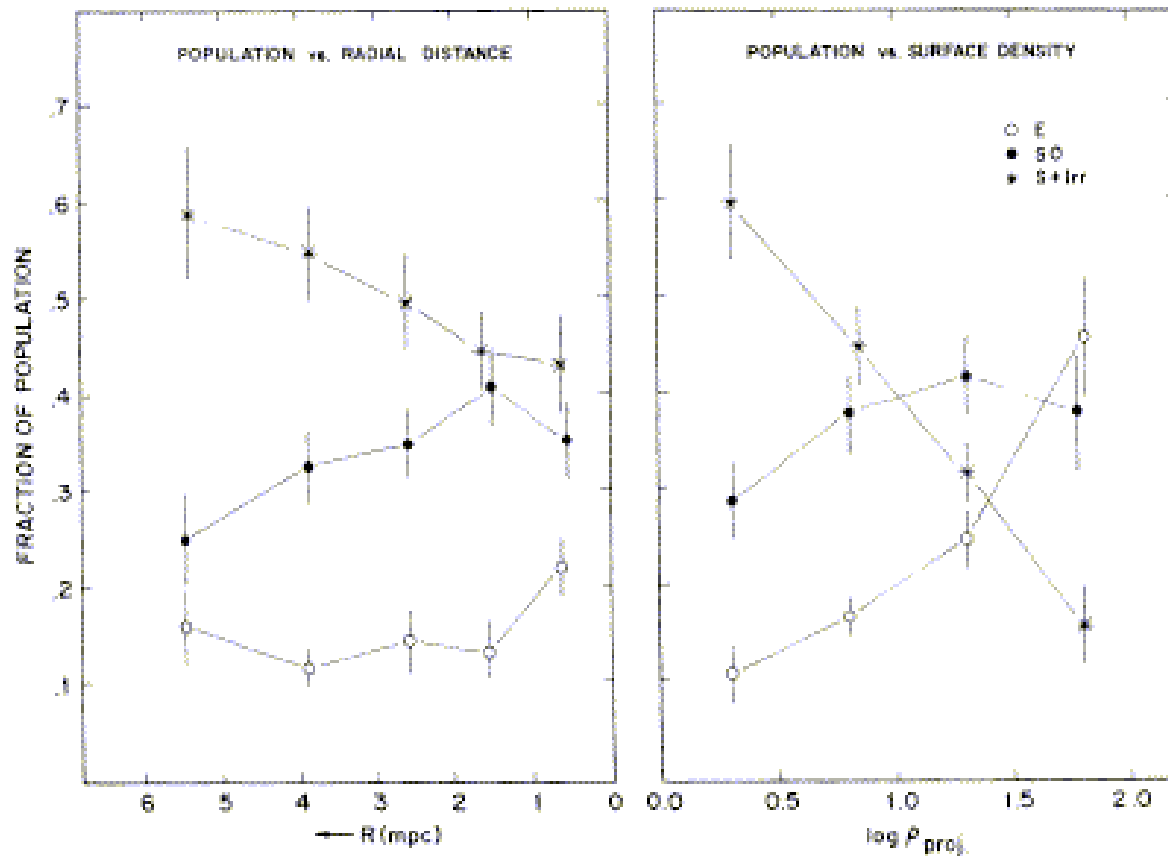
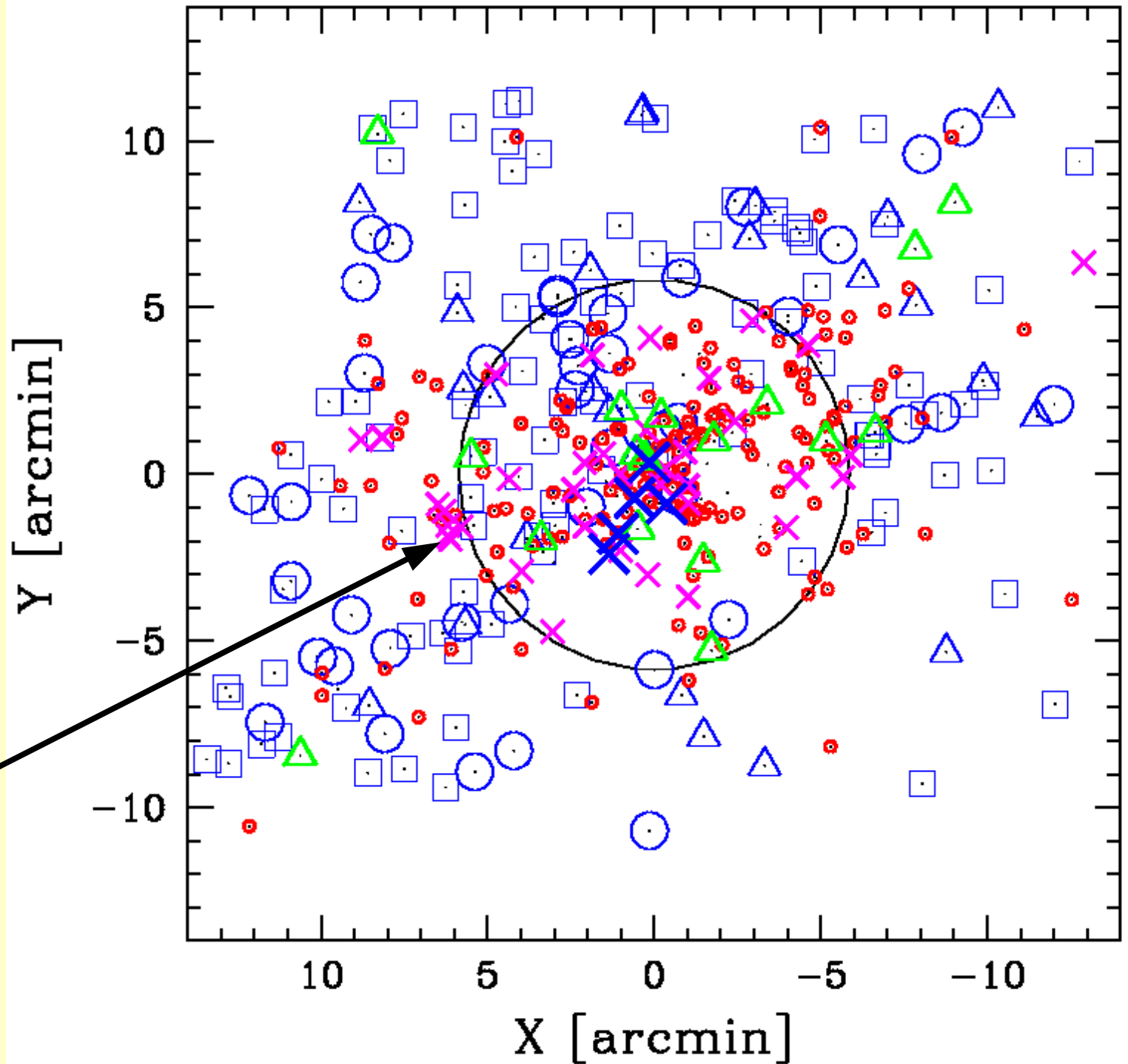


FIG. 5.—Population gradients in 6 moderately irregular clusters (A754, A993, A1736, A1983, 0326–53, 0559–40) as a function of radial distance from the cluster centroid and as a function of local surface density, showing the advantage of density as the free parameter.

Spatial Galaxy  
Distribution  
In cluster  
MACSJ1206

Girardi+2015



Different  
spectral types